

A wide-field H I study of the NGC 1566 group

Virginia A. Kilborn,^{1,2*} Bärbel S. Koribalski,² Duncan A. Forbes,¹ David G. Barnes³
and Ruth C. Musgrave¹

¹*Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Mail 31, PO Box 218, Hawthorn, VIC 3122, Australia*

²*Australia Telescope National Facility, CSIRO, PO Box 76, Epping, NSW 1710, Australia*

³*School of Physics, University of Melbourne, Parkville, VIC 3010, Australia*

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ABSTRACT

We report on neutral hydrogen observations of a $\sim 5.5 \times 5.5$ deg² field around the NGC 1566 galaxy group with the multibeam narrow-band system on the 64-m Parkes Telescope. We detected 13 H I sources in the field, including two galaxies not previously known to be members of the group, bringing the total number of confirmed galaxies in this group to 26. Each of the H I galaxies can be associated with an optically catalogued galaxy. No ‘intergalactic H I clouds’ were found to an H I mass limit of $\sim 3.5 \times 10^8 M_{\odot}$. We have estimated the expected H I content of the late-type galaxies in this group and find that the total detected H I is consistent with our expectations. However, while no global H I deficiency is inferred for this group, two galaxies exhibit individual H I deficiencies. Further observations are needed to determine the gas removal mechanisms in these galaxies.

Key words: surveys – galaxies: clusters: general – galaxies: individual: NGC 1549 – galaxies: individual: NGC 1553 – galaxies: individual: NGC 1566 – radio lines: galaxies.

1 INTRODUCTION

The majority of galaxies in the Universe reside in galaxy groups (e.g. Tully 1987). However, the physical processes operating in groups are poorly understood, and no detailed census of their contents is available. Studies of neutral hydrogen in groups have previously concentrated almost exclusively on Hickson compact groups (e.g. Shostak, Allen & Sullivan 1984; Verdes-Montenegro et al. 2001), as they require small observing regions. Studies of loose groups, which are more extended, are much fewer in number. The advent of the multibeam receiver on the Parkes Telescope (Staveley-Smith et al. 1996) has made the study of loose groups feasible, and the success of the instrument has been shown with new group members being found in H I surveys of optically well-studied groups. For example, 10 new members were found in the Centaurus A group as part of the H I Parkes All Sky Survey (HIPASS, Banks et al. 1999), and new dwarf galaxies were found in H I surveys of the NGC 5044 and 1052 groups (McKay et al. 2004). Several new galaxy groups were identified in the HIPASS Bright Galaxy Catalogue (BGC, Koribalski et al. 2004).

H I surveys of groups can reveal the effects of tidal interactions in these environments (Haynes, Giovanelli & Chincarini 1984). One of the best examples of this is the VLA survey of the M81 group by Yun, Ho & Lo (1994), where the H I distribution encompasses

the major galaxies in the group with many streams and tidal tails. Further tidal tails were observed when the M81 group was observed during the H I Jodrell All Sky Survey (HIJASS, Boyce et al. 2001). Blind H I surveys of groups can turn up unexpected results. For example, Barnes & Webster (2001) imaged the H I environment in five nearby galaxy groups, leading to the discovery of a ring of H I gas encircling the compact core of the loose group LGG 138. H I surveys can also uncover information about galaxy formation and evolution in groups. This was the case when a large bridge of H I was discovered with Australia Telescope Compact Array (ATCA) mapping of the NGC 6221/6215 galaxy group (Koribalski & Dickey 2004).

Cold dark matter (CDM) theory predicts many more low-mass galaxies surrounding large galaxies than are actually observed in groups (e.g. Moore et al. 1999; Klypin et al. 1999). Searches for these low-mass satellites have been made in H I in several nearby groups such as Centaurus A and Sculptor (de Blok et al. 2002), NGC 5798, 5962, 5970, 6278, 6500 and 6574 to a limit of $\sim 7 \times 10^6 M_{\odot}$ (Zwaan 2001), and NGC 1808 (Dahlem, Ehle & Ryder 2001). More recently, Pisano et al. (2004) surveyed three groups (LGG 93, 180 and 478), similar in size and structure to the Local Group, for low-mass H I clouds (to an H I mass sensitivity of $\sim 10^7 M_{\odot}$). No population of H I-rich satellites with little or no optical emission were uncovered in any of the above surveys, leading to the conclusion that massive H I clouds are rare in the group environment. However, to date, no H I surveys have been deep enough to rule out a population of low-mass ($< 10^7 M_{\odot}$) H I clouds.

*E-mail: vkilborn@astro.swin.edu.au

Pointed H I surveys of groups and clusters of galaxies have generally found a trend for galaxies to be H I deficient as a function of distance from their centre. Verdes-Montenegro et al. (2001) found in a homogeneous survey of 48 Hickson compact groups (HCGs) that they contained only 40 per cent of the expected H I mass, based on the optical properties of the individual galaxies. Solanes et al. (2001) analysed the H I properties of 18 clusters, and found that two-thirds of their sample were H I deficient in the central regions (within the Abell radius). However, a recent study of the H I content of compact groups of galaxies (Stevens et al. 2004) has found no strong evidence for an H I deficiency in their sample (15 groups). Their study highlights the large uncertainties inherent in determining the expected H I mass of a galaxy based on its optical properties.

We are conducting wide-field H I observations as part of the Group Evolution Multiwavelength Study (GEMS) (Osmond & Ponman 2004, hereafter OP04; McKay et al. 2004), containing ~ 60 groups selected to have existing *ROSAT* PSPC X-ray observations. OP04 detail the selection criteria and X-ray properties of the groups. Seventeen of these groups in the Southern hemisphere were surveyed for H I with the multibeam narrow-band system on the 64-m Parkes Telescope. The main aims of this H I survey are to provide a census of the H I gas in groups, to find new group members and possible intergalactic H I clouds, and to make a direct comparison between hot and cold gas in the group environment for the first time. Here we report on our H I results for the NGC 1566 group.

1.1 The NGC 1566 galaxy group

NGC 1566 is the brightest spiral member of a nearby group of galaxies in Dorado at a distance of about 21 Mpc (OP04). This distance and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ are used throughout this paper. The Dorado group (see Fig. 1) consists of at least 46 galaxies (Maia, Da Costa & Latham 1989) and covers a velocity range from ~ 500 to 2000 km s^{-1} . It is part of the Fornax Wall, which connects the NGC 1672, 1566 and 1433 galaxy groups.

Fig. 1 shows the H I distribution of the Dorado group. The positions of known optical galaxies are marked. This map was derived from the HIPASS¹ data cubes, which have a velocity resolution of 18 km s^{-1} (see Barnes et al. 2001). The AIPS task MOMNT was used with Hanning smoothing over five channels (65 km s^{-1}), five pixels (20 arcmin) and a flux density cut-off, after smoothing, of 20 mJy beam^{-1} per channel. The H I distribution of the region shows three distinct groupings of galaxies around the galaxies NGC 1433, 1672 and 1566. These three groups are at similar velocities and thus are part of the larger complex.

The NGC 1566 group has been the subject of several optical studies. Huchra & Geller (1982) catalogue 18 members in the group (which they call HG3). The group catalogue by Garcia (1993) lists six members, and they derive a group recession velocity of 1292 km s^{-1} and a velocity dispersion of 99 km s^{-1} . Ferguson & Sandage (1990), Morshidi-Esslinger, Davies & Smith (1999) and Carrasco et al. (2001) surveyed various areas of the Dorado group in the optical and catalogued large numbers of mostly low-luminosity dwarf galaxies as potential members. The LEDA data base lists 24 galaxies with known velocities that lie within the NGC 1566 group, while the NASA Extragalactic Database (NED) lists 22 previously catalogued galaxies in the region.

Several members of the NGC 1566 group have been previously mapped in H I at the ATCA. Walsh (2004) studied the H I dynamics of

the spiral galaxy NGC 1566 itself, finding a nearly circular H I envelope, and a total dynamical mass of $1.2 \times 10^{11} M_{\odot}$. Many other observations have been made of this object, including CO (Bajaja et al. 1995), H α (Pence, Taylor & Atherton 1990), H I (Reif et al. 1982) and X-ray and radio continuum (Ehle et al. 1996). NGC 1533 is a nearby, early-type galaxy with two small companions, IC 2038/9. Recent ATCA observations obtained by Ryan-Weber, Webster & Bekki (2003) reveal an asymmetric H I ring surrounding the optical galaxy, with only a small amount of H I associated with IC 2038 and none with IC 2039. The interacting pair of galaxies NGC 1596/1602 has recently been mapped in H I at the ATCA (Chung et al., in preparation).

OP04 studied the X-ray properties of the NGC 1566 group using *ROSAT* PSPC observations. They found galaxy halo emission around NGC 1566 itself ($\log L_X = 40.41 \text{ erg s}^{-1}$, see their table 4), but no group-scale emission. This suggests that the group is relatively young and not yet virialized. The NGC 1566 group centre and radius, as defined in the GEMS sample (see OP04, their table 1), is (J2000) $\alpha = 04^{\text{h}}20^{\text{m}}00^{\text{s}}.6$, $\delta = -54^{\circ}56'17''$ and $r_{500} = 0.47 \text{ Mpc}$, where r_{500} is the radius at 500 times the critical density of the Universe at the current epoch. OP04 derive a group velocity of $v = 1402 \pm 61 \text{ km s}^{-1}$ and a velocity dispersion of $\sigma_v = 184 \pm 47 \text{ km s}^{-1}$, based on the nine group members within $v \pm 3\sigma_v$ and r_{500} . The total B-band luminosity of the NGC 1566 group based on these nine members is $1.87 \times 10^{11} L_{\odot}$.

In this paper we present our H I survey of the NGC 1566 group. In Section 2 we summarize the observations and data reduction, along with the source detection and optical identification. In Section 3 we give the results from the H I survey. Discussion of the results is provided in Section 4, including the optical properties and H I content of the group, X-ray properties and dynamics. Finally we present our conclusions in Section 5.

2 OBSERVATIONS AND DATA REDUCTION

Our observations of the NGC 1566 group follow the approach described in McKay et al. (2004). We scanned a field of dimensions $5.5 \times 5.5 \text{ deg}^2$ using the Parkes 64-m Radiotelescope with the 21-cm Multibeam Receiver Staveley-Smith et al. (1996) installed at the prime focus. Observations were acquired in the periods 2002 January 11–18, 2002 March 8–9 and 2003 May 3–5. The field was centred at (J2000) $\alpha = 04^{\text{h}}16^{\text{m}}00^{\text{s}}$, $\delta = -55^{\circ}35'00''$. Scans were made at a rate of one degree per minute, along lines of equatorial latitude and longitude separated by 4 arcmin. A total of 198 scans were acquired.

Data were reduced using LIVEDATA, a component of the AIPS++ package. On input, raw correlator spectra were convolved with a Hanning filter to eliminate ringing from narrow and bright emission sources. Subsequent processing steps were based on those for HIPASS data Barnes et al. (2001) and previous GEMS data McKay et al. (2004), with changes made to improve dynamic range near bright sources, and to track the time evolution of the (source-free) sky at higher order. Bandpass correction was applied on a per-beam, per-scan basis by iteratively clipping the data and fitting a second-degree polynomial to the time series of each channel. Data acquired within 20 arcmin of known H I sources in the field (based on a first, quick process and search step) were masked during fitting to prevent contamination of the calculated bandpass. Each bandpass-removed spectrum was frame-shifted to the barycentre of the Earth–Sun system, and then baselined by subtracting a source-masked, iterative, clipped second-degree polynomial fit, this time in the spectral domain.

¹ <http://www.atnf.csiro.au/research/multibeam>

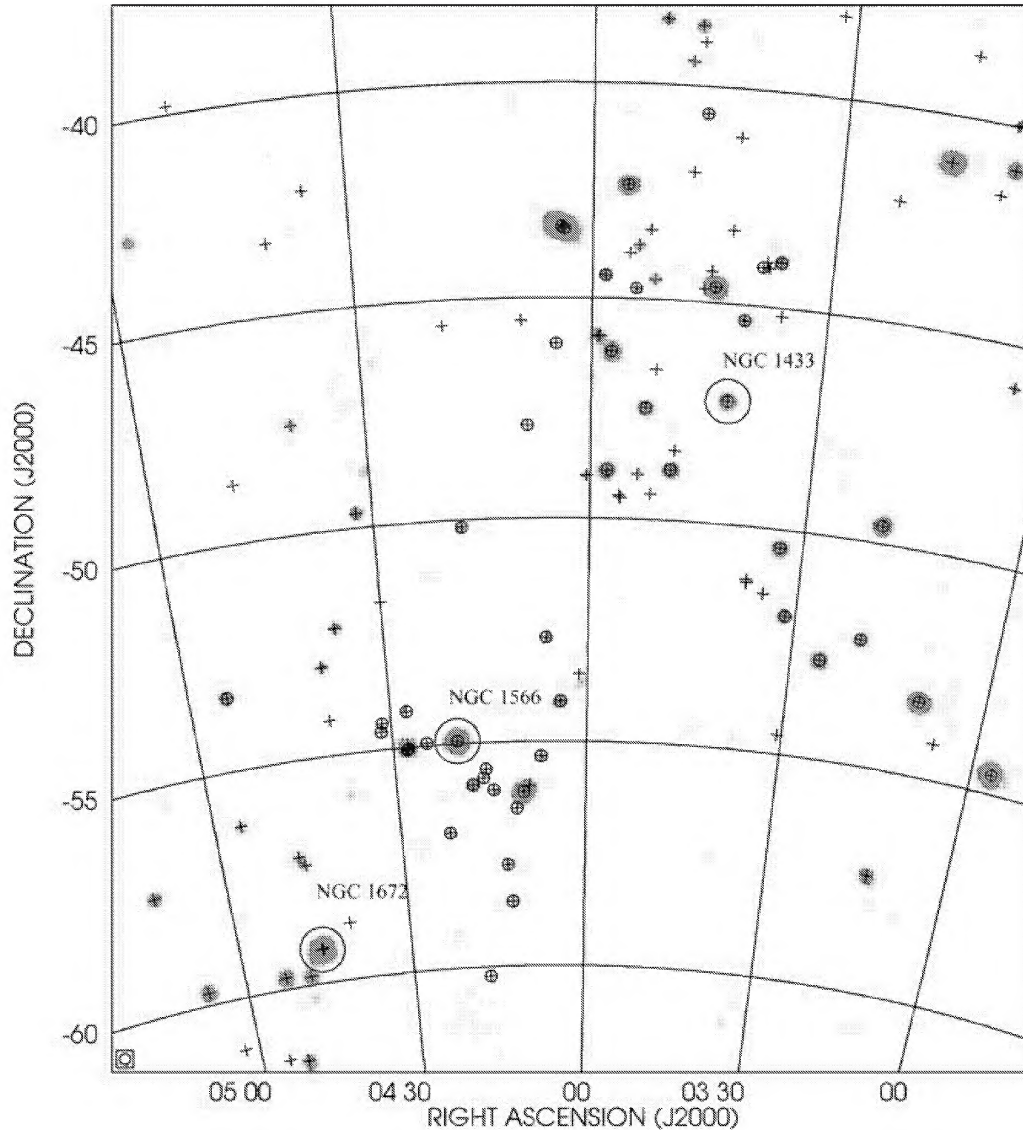


Figure 1. The $H\text{I}$ -rich galaxies in the Dorado group as revealed by HIPASS. The measured $H\text{I}$ flux densities are shown by the grey-scale. The gridded beam is 15.5 arcmin, and is displayed in the lower left corner. The crosses denote all galaxies in the velocity range from 500 to 2400 km s^{-1} recorded in NED, and the open circles denote the 46 members of the Dorado group as listed by Maia et al. (1989). Large circles mark the galaxies NGC 1433, 1566 and 1672, which dominate the smaller groups within the Dorado group.

Calibrated, processed spectra were imaged using the GRIDZILLA component of AIPS++, which is described in Barnes et al. (2001). We used the WGTMED statistic, which calculates mesh pixel values by taking the weighted median of data within 6 arcmin of the centre of each pixel.² The weight values were directly proportional to the canonical, Gaussian observing beam profile, which has a full width at half-maximum (FWHM) of 14.4 arcmin. We produced 2026 channel maps, centred on $V_{\text{CMB}} = 1231.8 \text{ km s}^{-1}$, with velocity widths of 0.824 km s^{-1} .

The gridded cube was smoothed in the spectral domain again with a Hanning filter three channels wide, and every second channel was discarded. The velocity resolution of the image was then measured,

yielding a FWHM spectral resolution of 2.6 km s^{-1} . The channel separation in the final cube is 1.65 km s^{-1} . The rms noise in the final cube is measured to be 18.8 mJy per channel. The final gridded beam size is 15.5 arcmin, and the pixel size is 4 arcmin square. The final cube parameters are summarized in Table 1.

Table 1. Narrow-band observing and cube parameters.

Gridded beam size	15.5 arcmin
Total observing time	18.5 h
Velocity range	400–2080 km s^{-1}
Channel width	1.65 km s^{-1}
Velocity resolution	2.6 km s^{-1}
rms noise per channel	18.8 mJy

² For comparison, HIPASS images use the GRIDZILLA MEDIAN statistic, which is the median of weighted values.

Table 2. H I properties of the detected galaxies.^a

No.	α, δ (J2000) (^h ^m ^s), ([°] ['] ^{''})	v (km s ⁻¹)	w_{50} (km s ⁻¹)	w_{20} (km s ⁻¹)	S_{peak} (Jy)	$F_{\text{H I}}$ (Jy km s ⁻¹)	rms (Jy)	Order	Box	$M_{\text{H I}}$ (10 ⁸ M _⊙)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
1	4:12:11, -58:34:17	1466±3	55±6	102±9	0.106±0.009	5.7±0.7	0.008	3	5	5.9±0.7
2	4:10:51, -56:30:32	1310±7	104±14	129±21	0.028±0.007	2.0±0.6	0.006	3	5	2.1±0.6
3	4:12:39, -57:45:50	1176±3	189±6	214±9	0.076±0.008	8.7±0.9	0.004	3	5	9.0±0.9
4	4:27:25, -57:27:21	1215±7	75±14	137±21	0.048±0.007	3.2±0.6	0.005	3	5	3.3±0.6
5	4:09:42, -56:06:55	796±3	247±6	324±9	0.382±0.020	73.7±2.9	0.012	3	13	76.5±3.0
6	4:22:42, -56:16:59	1350±4	52±8	76±12	0.049±0.007	2.4±0.5	0.005	7	5	2.5±0.6
7	4:14:40, -56:04:25	1298±5	189±10	343±15	0.140±0.010	23.2±1.3	0.011	3	13	24.1±1.4
8	4:17:56, -55:57:06	1370±2	193±4	218±6	0.124±0.009	16.4±1.1	0.004	3	5	17.1±1.2
9	4:07:10, -55:18:01	1068±3	99±6	119±9	0.073±0.008	5.9±0.7	0.005	3	5	6.1±0.8
10	4:19:53, -54:56:46	1502±1	200±2	223±3	1.040±0.052	148.1±6.4	0.009	3	13	153.9±6.7
11	4:27:45, -55:01:08	1572±3	93±6	174±9	0.283±0.016	28.1±1.6	0.009	3	11	29.2±1.7
12	4:03:56, -54:05:04	1180±2	357±4	381±6	0.101±0.009	15.9±1.1	0.004	3	5	16.5±1.2
13	4:05:42, -52:41:01	902±2	88±4	103±6	0.079±0.008	4.8±0.6	0.007	3	5	5.0±0.7

^aThe columns are as follows: (1) GEMS galaxy number. (2) Fitted H I centre position. (3) Heliocentric velocity in the optical convention. (4) The 50 per cent velocity width. (5) The 20 per cent velocity width. (6) H I peak flux density. (7) Integrated H I flux density. (8) Clipped rms per channel. (9) Order of baseline fit to the H I spectrum. (10) Box size used to create the H I spectrum in 4-arcmin pixel. (11) H I mass for the detection assuming cluster distance of 21 Mpc. The errors are derived following Koribalski et al. (2004).

2.1 Source finding

The H I data cube was searched visually for sources using the KVIEW visualization program (Gooch 1995). Two of us (VAK and RCM) catalogued sources from the original cube, and also from two smoothed versions of the cube, with final velocity resolution of 3.3 and 6.6 km s⁻¹. The rms in the smoothed cubes was 11 and ~8 mJy beam⁻¹ respectively. To determine the detection limit for sources in the cube, and as an extra check for the searching, 20 fake point sources were injected randomly into the original cube. These sources had a range of peak flux density between 10 and 30 mJy, and velocity width of 50–500 km s⁻¹. Two fake-injected cubes were made, to ensure that, if one of the randomly placed fake sources lay upon a real H I signal, the other cube would likely have it unobscured. Both the unsmoothed and two smoothed versions were searched for sources, as with the original data cube. From the recovery rate of fake sources, we have determined our survey to be complete to a peak flux density limit of 58 mJy ($\approx 3\sigma$) in the unsmoothed cube. While we did detect a number of sources below this limit, the sources with the lowest peak flux density were detected in the smoothed cubes only. No relationship with velocity width was seen with the detectability of these fake sources. Assuming a galaxy velocity width of 100 km s⁻¹ and a Gaussian H I profile, this translates to an H I mass limit of $3.5 \times 10^8 M_{\odot}$ ($\sim 1 \times 10^9 M_{\odot}$ for a velocity width of 300 km s⁻¹).

2.2 Derivation of the H I parameters

Once the source list was made, the H I parameters for each detection were derived using MIRIAD routines (Sault, Teuben & Wright 1995) on the smoothed cube (velocity resolution 6.6 km s⁻¹). A zeroth-order moment map (MOMENT) was made for the velocity range of emission of each detection. The central point of emission was then determined by using IMFIT to fit a two-dimensional Gaussian to this moment map. This central position was then used in the task MBSPECT to make a spectrum of the source.

A spectrum for each source was produced using a box size of 5 pixels in width (20 arcmin), assuming a point source. The 20 and 50 per cent velocity widths (taken as the maximum width fit in MBSPECT), integrated flux density (robust moment 0) and peak flux

density for each source were also determined with MBSPECT. The heliocentric velocity for a source was determined to be the centre point of the 20 per cent velocity width (in the optical convention).

The integrated flux density was also measured for a box size of 7 pixels to determine if any extended emission surrounded the detection. Four of the galaxies were found to have extended emission. These four sources were fitted in a similar way to above, but in this case the flux within the box was summed rather than assuming a point source, and the box size was iteratively increased until the total flux density was constant. The final parameters for all galaxies detected in H I are given in Table 2, including the fitting parameters for the spectra. The H I mass of each detection was determined using $M_{\text{H I}} = 2.356 \times 10^5 D^2 F_{\text{H I}}$, where $F_{\text{H I}}$ is the integrated H I flux density in Jy km s⁻¹ (see Table 2), and D is the distance to the group in Mpc (21 Mpc for the NGC 1566 group).

H I spectra for each of the galaxies are shown in Fig. 2. These spectra show the fitted baseline and the region excluded from the baseline fit with the vertical lines; the 20 and 50 per cent velocity widths are shown by the crosses (minimum) and circles (maximum); the peak flux density is shown with a filled circle.

Uncertainties on the derived parameters were calculated according to Koribalski et al. (2004). Detections with low H I flux have large uncertainties in the determined velocity, as do detections with asymmetric profiles. The uncertainty on the peak flux density is always greater than the rms noise of the fitted cube. Positional uncertainties are not quoted. However, the positional uncertainty of the H I detections can be calculated as the gridded beam size divided by the signal-to-noise ratio (Koribalski et al. 2004).

2.3 Optical identification of the H I detections

Optical identification of the H I detections was made using the NASA Extragalactic Database (NED). A 6-arcmin region around the central point of the H I emission was searched for previously catalogued galaxies of either the same redshift or unknown redshift in NED. Table 3 gives the identifications for the H I detections, with their previously determined velocities. Eleven H I sources in the NGC 1566 group had matching optical counterparts. Of these, two H I sources (5 and 11) correspond to catalogued galaxy pairs. The remaining two H I detections had no previously catalogued galaxy

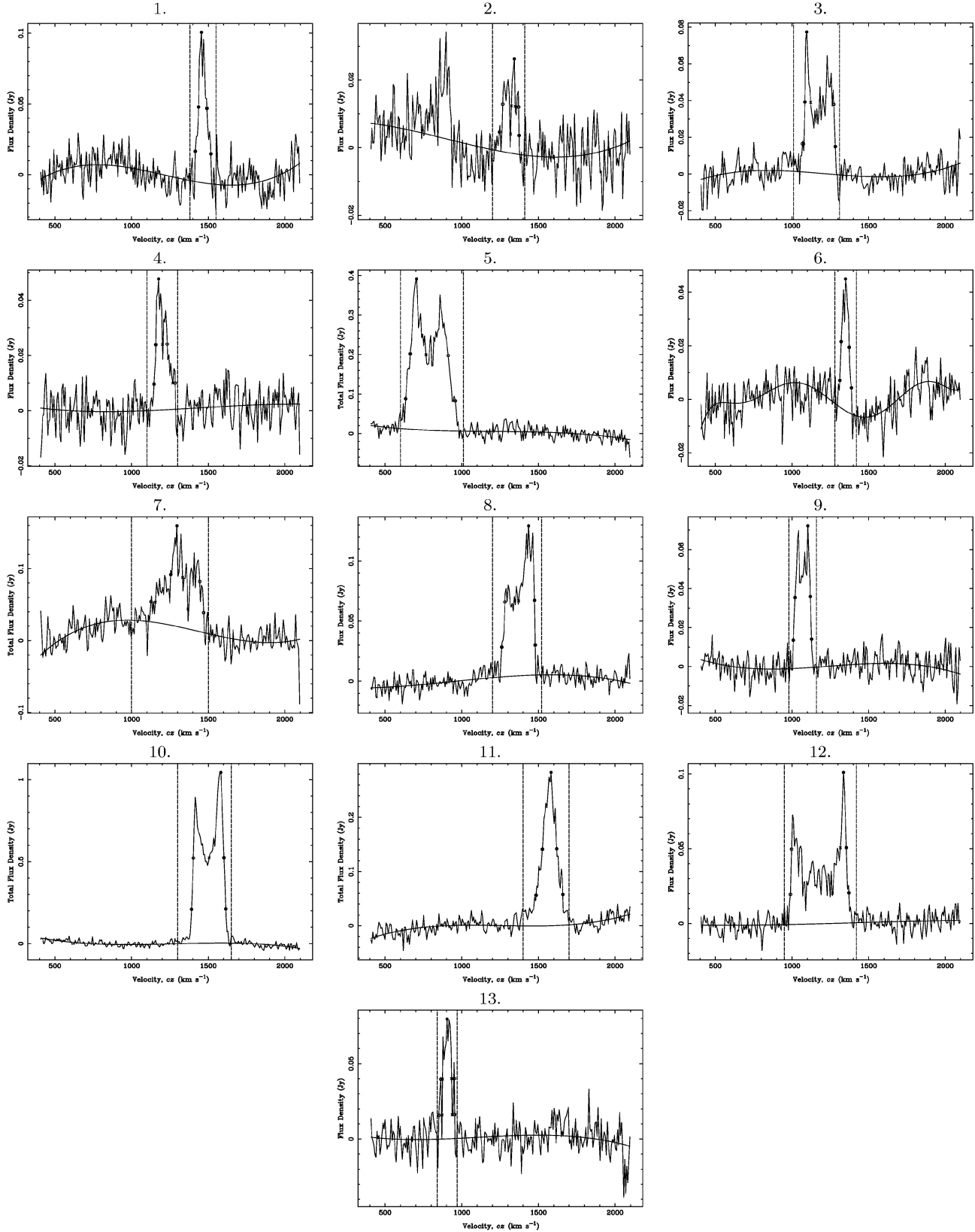


Figure 2. H_1 spectra of the 13 galaxies detected in our NGC 1566 survey. The spectra were Hanning smoothed to a velocity resolution of 6.6 km s^{-1} . The fitted baseline is shown and the H_1 peak flux density is marked with a filled circle. The w_{20} and w_{50} velocity widths are shown by the open circles (outer fit) and crosses (inner fit). The velocity region between the vertical lines in the spectra was disregarded in the baseline fit.

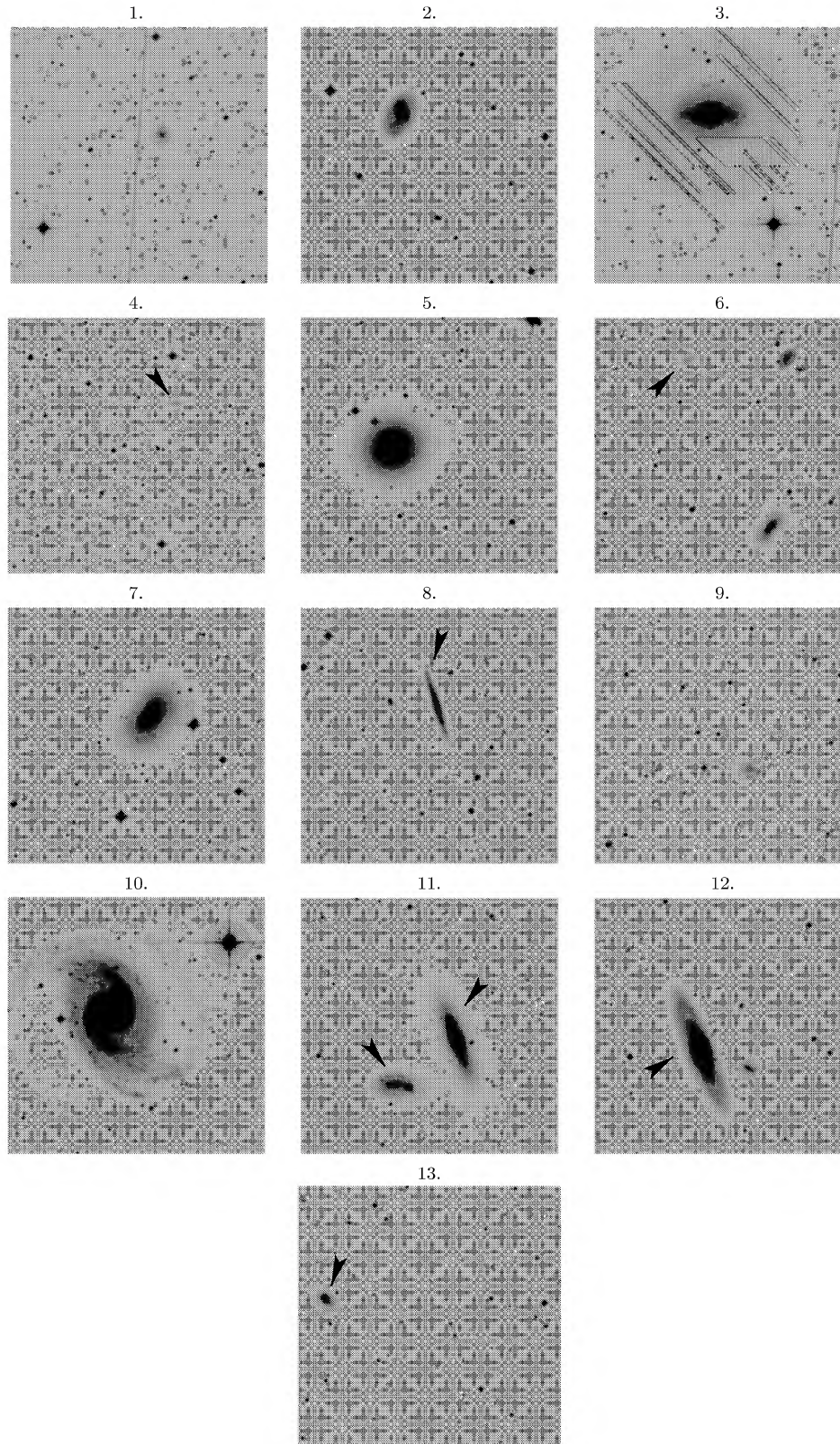


Figure 3. Second-generation DSS *R*-band images of the 13 galaxies detected in the NGC 1566 group. Each image is 10 arcmin across, and is centred at the H I detection. When there is more than one galaxy in the field of view, arrows indicate the matching optical counterpart, determined by either the known optical velocity of the galaxy, or in the case of images 4 and 6, from high-resolution ATCA observations. Note that in fields 8 and 12 there are several galaxies within the field of unknown redshift.

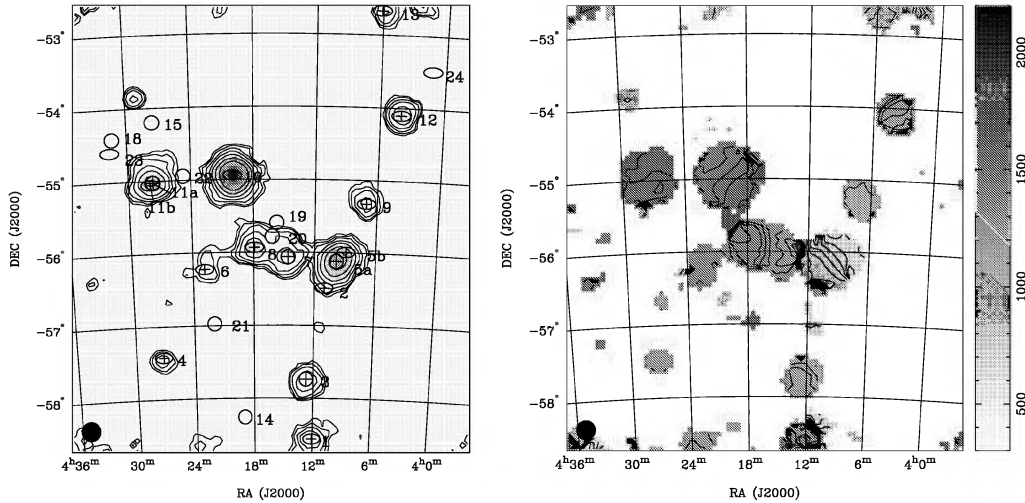


Figure 4. (a) (left) The integrated H I intensity map of the NGC 1566 group as obtained from the H I Parkes narrow-band observations. The measured H I flux densities (contours) are 0.5, 1, 2, 4, 8, 16 and 32 Jy beam⁻¹ km s⁻¹. Overplotted are the positions of known group members. The 13 crossed symbols are those detected in our H I survey, and the open symbols were undetected in H I. The circles represent E/S0 galaxies, and ellipses represent late-type galaxies. The numbers correspond with those in Table 3. Note that galaxies 16 and 17 have been previously associated with this group, but lie outside our survey region. (b) (right) The velocity distribution for the NGC 1566 group. The contour levels are 600 to 1615 km s⁻¹ in steps of 35 km s⁻¹. The Parkes beam is shown in the bottom left corner of each image.

of known redshift nearby, and further high-resolution observations were necessary to determine the optical counterparts to these sources (see Section 3 for details). Optical images for each detection were downloaded from the Digital Sky Survey³ for each of the galaxies. These can be seen in Fig. 3. The DSS images are 10 arcmin square, and are in *R* band. A variety of galaxy types can be seen, from bright spirals and elliptical galaxies to small, faint dwarf galaxies. The optical characteristics for all galaxies are given in Table 3.

3 RESULTS

3.1 H I properties of the NGC 1566 group

The H I distribution for the NGC 1566 group is shown in Fig. 4(a), and the mean H I velocity field is shown in Fig. 4(b). These were derived from the narrow-band data cubes. After Hanning smoothing the data to the same spectral resolution as HIPASS, we used the AIPS task MOMNT (same smoothing parameters as for Fig. 1) to derive these H I moment maps. Because of varying baseline curvature (see Fig. 2), these moment maps are not as sensitive as the individual H I spectra. We used a flux density cut-off of 10 mJy beam⁻¹.

We detected 13 sources in the H I data cube, compared with 24 optical galaxies with known velocities in the region from LEDA. Two of the H I sources correspond to interacting galaxy pairs. We also determined the redshift for two galaxies previously not known to be group members, thus taking the number of known members for this group to 26. The spatial distribution of all known galaxies in the NGC 1566 group is overplotted on the H I distribution in Fig. 4(a). The crossed symbols are those detected in H I and open symbols were not detected in our survey. The ovals represent late-type galaxies, and the circles are E/S0 galaxies. The H I distribution shows an apparent connected feature encompassing galaxies 2, 5a,

5b, 6, 7, 8 and 20. However, except for around the interacting system of NGC 1533/IC 2038, it is likely that this apparent feature is due to smearing of the beam in the H I data set, and higher-resolution H I observations of the region are needed to determine the nature of the H I distribution.

Nine known members of the NGC 1566 group were not detected in our H I survey (to an H I mass limit of $\sim 3.5 \times 10^8 M_{\odot}$), and two further members were outside the survey region. We detected four E/S0 galaxies in H I, but did not detect seven other early-type galaxies previously determined to be in the group. There was one irregular and one Sa galaxy not detected in H I, but all other known spiral/late-type galaxies were detected in H I. NGC 1553 has previously been reported as detected in H I (de Vaucouleurs et al. 1991); however, the resolution of our Parkes data is too coarse to determine a separate detection for this galaxy.

The two new group members were observed at the ATCA from 2004 March 3–7. The data were reduced using standard MIRIAD routines, and the resulting H I distributions are shown in Fig. 5. Both of the H I detections correspond to previously catalogued low surface brightness irregular galaxies with blue luminosities around $10^8 L_{\odot}$. These two galaxies are LSBG F157–081 and APMBGC 157+016+068 (as named in NED), and both have no previous known redshift. The H I mass (as determined from the Parkes data) for the two galaxies is $3.3 \times 10^8 M_{\odot}$ and $2.5 \times 10^8 M_{\odot}$, respectively. Both of these new group members were unresolved in the ATCA H I maps.

The H I mass of the sources we detected in the Parkes data range from 2.1×10^8 to $1.4 \times 10^{10} M_{\odot}$ (see Fig. 6). The highest H I mass galaxy was NGC 1566 itself. This galaxy is extended in the Parkes beam. We detect an integrated H I flux density of 148.1 ± 6.4 Jy km s⁻¹, which is consistent with previously measured values. The HIPASS Bright Galaxy Catalogue (BGC, Koribalski et al. 2004) measured 140 ± 10 Jy km s⁻¹, as do Mathewson & Ford (1996). Walsh (2004) measured 147 ± 10 and 161 ± 15 Jy km s⁻¹ from Parkes and the ATCA, respectively, for NGC 1566. The total H I mass we detected in the NGC 1566 group is $3.5 \times 10^{10} M_{\odot}$; thus

³ The Digital Sky Survey, provided by the Space Science Institute, based on photographic data from the UK Schmidt Telescope.

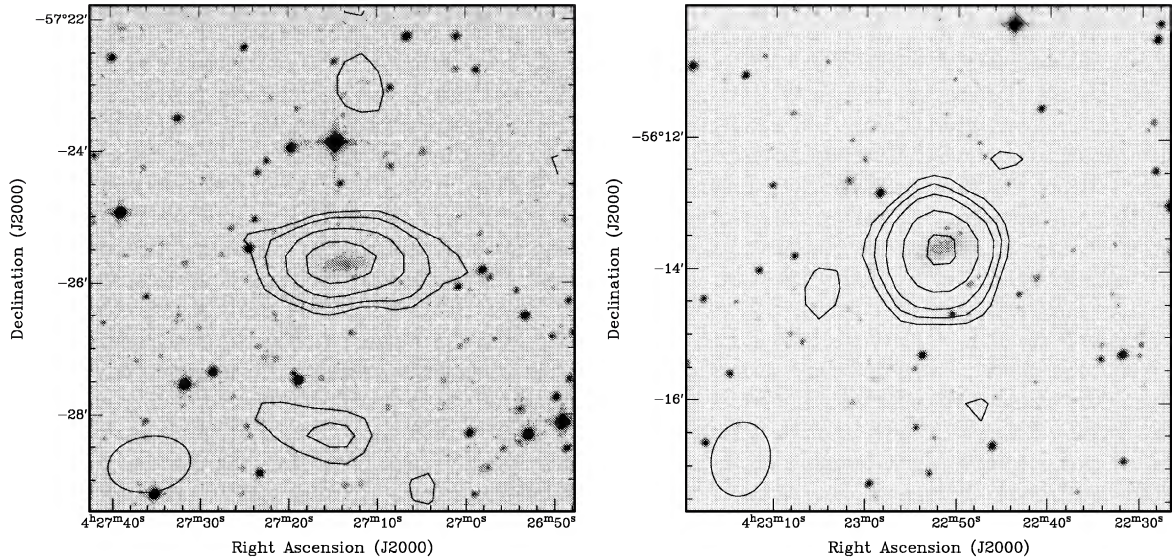


Figure 5. ATCA H I distribution overlaid on DSS II *R*-band images of the two new group members of the NGC 1566 group, namely LSBG F157–081 (left) and APMBGC 157+016+068 (right). The contour levels are 0.15, 0.25, 0.4, 0.75, 1.25 Jy km s⁻¹ (starting at 0.25 Jy beam⁻¹ km s⁻¹ for the left image). The ATCA beam is shown in the bottom left corner of each image.

the galaxy NGC 1566 itself accounts for nearly half the H I in the group. We found that over half the galaxies in this group have H I masses less than $10^9 M_{\odot}$.

3.2 Optical properties of the NGC 1566 group

The total *B*-band magnitudes for all known members in the NGC 1566 group are listed in Table 3. Total luminosities were determined from these magnitudes [after the magnitudes were corrected for Galactic extinction (Schlegel, Finkbeiner & Davis 1998)], and are also listed in this table. The highest-luminosity galaxy in the NGC 1566 group is NGC 1553, with a luminosity of $4 \times 10^{10} L_{\odot}$. NGC 1553 is an S0 galaxy that was undetected in our H I survey. The second brightest galaxy in the group is NGC 1566 with $L_B = 3.7 \times 10^{10} L_{\odot}$. The two lowest-luminosity galaxies are the two new members detected in our H I survey, LSBG F157–081 and APMBGC 157+016+068, with luminosities of 9×10^7 and $1.3 \times 10^8 L_{\odot}$, respectively.

The H I mass to *B*-band luminosity ratio (M_{HI}/L_B) for the galaxies in the NGC 1566 group varies from <0.01 for the early-type galaxies NGC 1553 and 1549, to 3.7 for the dwarf irregular galaxy LSBG F157–081. The total *B*-band luminosity for the group is $1.8 \times 10^{11} L_{\odot}$, compared to the total H I mass of the group of $3.5 \times 10^{10} M_{\odot}$. This gives an overall M_{HI}/L_B of 0.19 M_{\odot}/L_{\odot} . The two new group members have M_{HI}/L_B of 3.7 and 1.9, which confirms that H I surveys are a good way of detecting gas-rich, low-luminosity group members.

Using the *B*-band magnitudes, we have constructed a luminosity histogram for this group. This is given in Fig. 6, which shows the total luminosity distribution for the group, and that for only those galaxies detected in H I (note that only single H I detections are included in the second sample). We find that the H I detected galaxies span the full range of the optical luminosity distribution, and have a similar shape to the total luminosity distribution. Similar to the H I mass distribution for this group, half of the galaxies in the group have $L_B < 10^9 L_{\odot}$. The optical galaxy numbers drop off towards L_B of

$10^8 L_{\odot}$, which corresponds to the optical limit for previous redshift surveys of the region.

4 DISCUSSION

4.1 H I content and deficiency

H I observations of galaxies in groups and clusters have shown some galaxies to be H I deficient near the centre of the group or cluster (Giovannelli & Haynes 1985; Solanes et al. 2001; Verdes-Montenegro et al. 2001). In order to determine whether the H I content of a galaxy in a group or cluster is deficient, normal or overabundant, a good estimate of the expected H I is needed. The expected H I mass of isolated galaxies has been studied in the past using either the optical morphology and optical diameter of the galaxies (Haynes & Giovanelli 1984, hereafter HG84; Chamaraux et al. 1986; Solanes et al. 1996) or using blue magnitude and optical morphology (HG84). HG84 show that the H I content is more closely tied to the galaxy’s optical extent and morphological type; thus we use this method in the following analysis.

Optical properties for most previously catalogued galaxies in the HIPASS BGC (Koribalski et al., in preparation) were obtained from LEDA, including morphological type, optical diameters and magnitudes (Koribalski et al. 2004). Following Solanes et al. (1996), we made a linear regression fit for optical diameter and H I mass to nearly 800 BGC galaxies. The galaxies were divided into morphological type from Sb to Irr, and the resulting linear regression coefficients, α and β (such that $M_{\text{HI,ex}} = \alpha + \beta \log d$, where d is the optical diameter of the galaxy in kpc) for each morphological type are listed in Table 4. The scatter in the H I masses, σ , and the number galaxies at each morphological type, N , are also given in Table 4. We use the morphology and optical diameter of the NGC 1566 group galaxies, and the coefficients given in Table 4, to determine their expected H I mass, $M_{\text{HI,ex}}$. There are very small numbers of early-type galaxies in the BGC, and thus there is a large error in calculating the expected H I mass for galaxies of optical morphology earlier

Table 3. Group members of the NGC 1566 group. Galaxies 1–13 were detected in the H I survey, and galaxies 14–24 have optical detections only.^a

Galaxy No. (1)	Galaxy name (2)	Velocity (3)	Morphology (4)	B-band magnitude (5)	L_B (6)	M_{HI}/L_B (7)
1	IC 2049 ^b	1086±46	Sd	14.56	6.6	0.9
2	NGC 1536	1565±104	Sc pec	13.29	21.3	0.1
3	NGC 1543	1094±32	S0	10.67	237.9	0.04
4	LSBG F157-081	–	Irr	16.72	0.9	3.7
5a	NGC 1533	668±41	S0	11.74	88.8	–
5b	IC 2038	712±52	Sd pec	14.82	5.2	–
6	APMBGC 157+016+068	–	Irr	16.32	1.3	1.9
7	NGC 1546	1160±66	S0	12.31	52.5	0.5
8	IC 2058	1268±88	Sc	13.90	12.1	1.4
9	IC 2032	1070±34	Im	14.78	5.4	1.1
10	NGC 1566	1449±19	Sbc	10.19	370.2	0.4
11a	NGC 1596	1465±24	S0	12.01	69.2	–
11b	NGC 1602	1731±31	Im	13.79	13.4	–
12	NGC 1515	1216±48	Sbc	11.96	68.5	0.2
13	NGC 1522	1012±91	pec	14.08	10.3	0.5
14	ESO 118-019 ^c	1133±77	S0 pec	14.90	4.8	<0.7
15	ESO 157-030	1341±51	E4	14.30	8.4	<0.5
16	ESO 157-047 ^{d,e}	1733±39	S0/a pec	15.52	2.7	<1.29
17	ESO 157-049 ^d	1729±34	Sbc	14.37	7.9	<0.4
18	IC 2085	1010±30	S0 pec	14.26	8.7	<0.4
19	NGC 1549	1214±21	E0-1	10.48	283.4	<0.01
20	NGC 1553 ^f	1239±18	S0	10.10	402.2	<0.01
21	NGC 1574	925±66	S0	11.17	150.1	<0.02
22	NGC 1581	1527±88	S0	13.36	20.0	<0.18
23	NGC 1617	1063±21	Sa	11.35	127.2	<0.03
24	PGC 429411	1135±40	Irr	16.19	1.5	<2.3

^aThe columns are as follows: (1) GEMS galaxy number. (2) Galaxy name. (3) Optical velocity (km s^{-1}) from the RC3 catalogue (de Vaucouleurs et al. 1991), except where listed in the table notes below. (4) Morphological classification. (5) B-band magnitude. (6) B-band luminosity in units of $10^8 L_{\odot}$. (7) M_{HI}/L_B in units of M_{\odot}/L_{\odot} . No value is given for the interacting galaxies, and the upper limits are based on the H I mass limit of the survey of $3.5 \times 10^8 M_{\odot}$. The B-band magnitudes come from the ESO-LV catalogue (Lauberts & Valentijn 1989), apart from galaxies 4 and 6, which come from Maddox et al. (1990), and galaxy 24, which is the extinction-corrected total magnitude from the LEDA data base.

^bThe H I velocity for this galaxy is $1466 \pm 3 \text{ km s}^{-1}$, which is inconsistent with the published optical value.

^cVelocity from Saunders, Sutherland & Maddox (2000).

^dThese galaxies have previously been associated with the NGC 1566 group but lie outside our H I survey region.

^eVelocity from Loveday et al. (1996).

^fVelocity from Katgert et al. (1998).

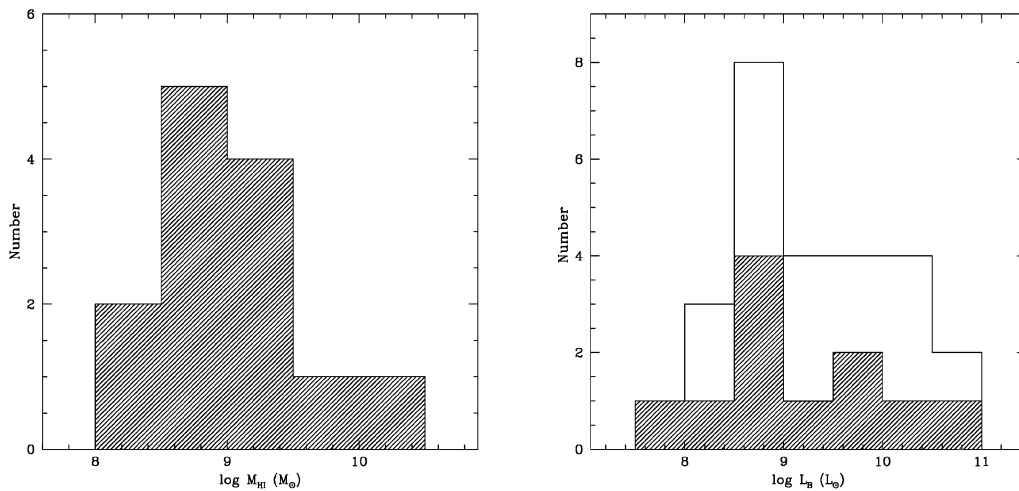
**Figure 6.** (left) H I mass distribution for NGC 1566 group. (right) Luminosity distribution for known members of the group. The open histogram indicates all members known optically, and the filled histogram indicates those galaxies detected in H I.

Table 4. Linear regression coefficients as derived from the HIPASS BGC (Koribalski et al. 2004).

Type	α	β	σ	N
Sb	8.39	1.02	0.24	85
Sbc	8.39	1.05	0.29	130
Sc	8.30	0.99	0.27	134
Scd	7.80	1.38	0.24	88
Sd	7.12	1.97	0.26	75
Sm	7.70	1.45	0.28	84
Irr	7.29	1.90	0.33	141

than Sb. In particular, it is known that elliptical galaxies are rarely detected in H I (Sadler 2001), and we do not discuss these galaxies further in this section.

We believe that the H I estimates from the BGC for late-type galaxies, especially Sd–Irr, are more accurate than previous estimates. Low surface brightness late-type galaxies can be easily missed in optical surveys, and thus these potentially H I-rich galaxies will not have been included in a targeted H I survey. As late-type galaxies have a high detection rate in H I, we expect to have a more complete sample than optically based pointed H I surveys. Indeed, the latest H I estimate for late-type galaxies in the literature was HG84, who have a sample size of just seven for irregular galaxies, whereas there are 141 in the BGC. We must bear in mind that we have no indication of how many late-type galaxies were not detected in the survey; thus the upper limit to the H I mass of undetected galaxies is not included in the linear regression fit.

We have calculated the expected H I mass for each galaxy in the NGC 1566 group, including previously catalogued members of known redshift that were not detected in our H I survey. To reduce the uncertainty in the calculation of expected H I mass, we limit our comparison to non-interacting galaxies of type Sb or later. Nine galaxies meet this criterion, including eight galaxies detected in our H I survey, and the undetected irregular galaxy PGC 429411. For comparison, we calculate the expected H I mass using both the coefficients in Table 4, and those given in HG84. Fig. 7 shows the

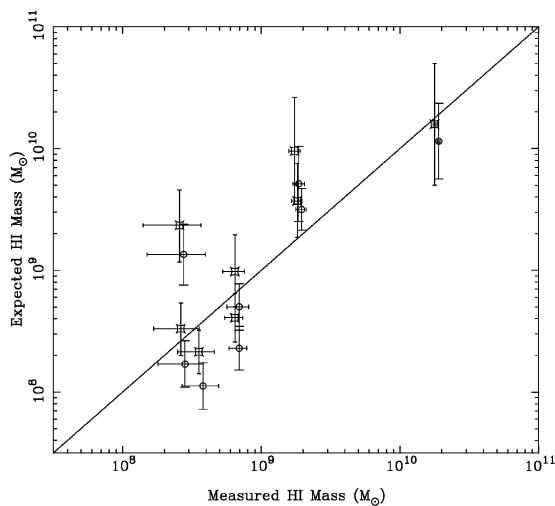


Figure 7. Expected H I mass of non-interacting late-type galaxies compared to the detected H I mass. The squares show the expected H I mass using the coefficients in Table 4, and the circles are the expected H I mass from the coefficients in HG84. The errors show the scatter in H I mass for a galaxy of given optical diameter and morphology.

expected H I mass versus the detected H I mass for the eight galaxies detected in our survey. The vertical error bars represent the scatter in the expected H I mass of a given H I mass and optical diameter for the BGC and HG84 samples. This figure shows that the expected H I mass from HG84 is always lower than that from the BGC.

The expected H I mass for most of the galaxies is consistent with what was observed, and, considering the group as a whole, the NGC 1566 group does not appear to be H I deficient. Two galaxies have a slightly higher detected H I mass than was expected, and a further two galaxies display a marked H I deficiency. These latter galaxies are the Sbc spiral NGC 1515 and the peculiar Sc spiral NGC 1536, with 5–10 times less H I detected than expected. While NGC 1515 is relatively isolated, NGC 1536 is situated spatially near the interacting galaxy pair of NGC 1533/IC 2038. However, there is a large velocity difference between the pair and NGC 1536 of nearly 500 km s^{-1} . The expected H I mass for the undetected irregular galaxy PGC 429411 is $3.1 \pm 1 \times 10^8 M_{\odot}$. As this H I mass lies near our detection limit for the H I survey, we are unable to tell whether this galaxy is H I deficient, or contains H I within the normal range.

4.2 Group distribution and dynamics

To study the dynamical state of the NGC 1566 group, we now derive the virial radius and velocity dispersion of the group members. The virial radius can be calculated using X-ray emission, where the latter exists in the intra-group medium. The *ROSAT* PSPC data for the NGC 1566 group, along with other GEMS groups, are presented in OP04. The X-ray data cover a circular region of ~ 2 deg in diameter. OP04 find X-ray emission that is consistent with coming from NGC 1566 itself, i.e. no evidence for extended emission from the intra-group environment.

OP04 derived an r_{500} radius of 470 kpc based on the X-ray temperature of $0.70 \pm 0.11 \text{ keV}$. However, the X-ray emission in the NGC 1566 group is dominated by hot gas in the halo of NGC 1566 itself, which reaches the background level at a radius of 29 kpc. So the temperature-based r_{500} should be treated with some caution.

An alternative method for calculating r_{500} is to use the group's velocity dispersion and the virial theorem (see OP04). The accuracy of this method improves with the number of reliable galaxy velocities. Based on only nine galaxies, OP04 derive a velocity dispersion for the group of $184 \pm 47 \text{ km s}^{-1}$. This would translate into an r_{500} radius of 250 kpc.

A better estimate of the dispersion for the group uses all known group members. Velocities were taken from the literature for the optically known galaxies not detected in our H I survey. Without removing any spatial or velocity outliers, we use a biweight estimator with bootstrap errors as per Beers, Flynn & Gebhardt (1990). This gives the average velocity of the group as $v = 1250 \pm 57 \text{ km s}^{-1}$, and velocity dispersion of $\sigma_v = 282 \pm 30 \text{ km s}^{-1}$. This velocity dispersion gives a r_{500} radius of 386 kpc, and thus a virial radius, r_v , of $\sim 580 \text{ kpc}$ for the NGC 1566 group ($r_v \sim 3/2 \times r_{500}$).

Fig. 8 shows the velocity distribution versus the distance from the centre of the group. The solid line and dashed lines show v and σ_v respectively. For this plot, the galaxy 19, NGC 1549, was chosen to be the centre position, as it is the brightest elliptical in the NGC 1566 group, and has a velocity very close to the mean velocity of the group.

Eleven galaxies associated with the NGC 1566 group lie within the virial radius of 580 kpc. A further 12 galaxies are located at $1\text{--}2 r_v$. As such a large number of galaxies lie beyond the virial radius, it is possible that this group is young and not yet virialized. There has been recent work on the evolution of galaxies lying outside

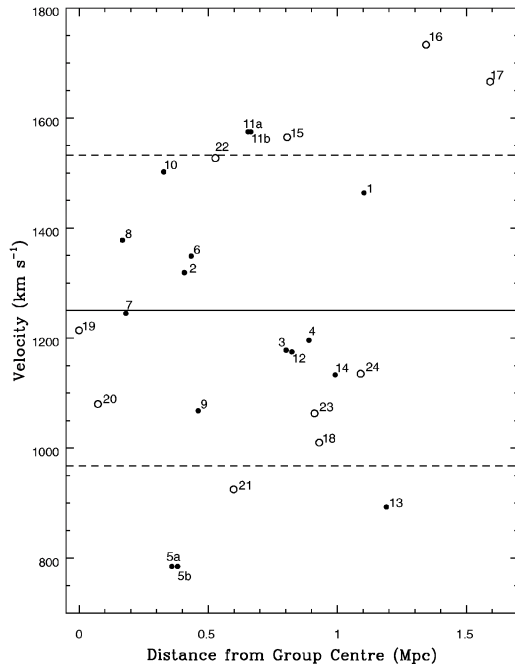


Figure 8. Velocity of the NGC 1566 group members versus their distance from the group centre. The filled circles show galaxies detected in H I, and the open symbols are known galaxies in the group that we did not detect. The solid line shows the mean velocity for the group, and the dashed lines show the velocity dispersion. The virial radius for this group is 580 kpc. Galaxy 19 (NGC 1549) is taken as the centre of the group.

the virial radius of groups and clusters. Simulations have shown that up to half of the galaxies that lie beyond the virial radius of a cluster may have travelled through the cluster centre (Ghigna et al. 1998; Gill, Knebe & Gibson 2004). Thus the evolutionary history of these ‘backsplash’ galaxies may have been influenced by this encounter, and mass loss of up to 40 per cent or greater might be expected. Observationally, there is some evidence of H I stripped galaxies not only in the centres of clusters (e.g. Solanes et al. 2001), but also nearer the edges of clusters (Vollmer 2003; Kenney, van Gorkom & Vollmer 2004; Vogt et al. 2004). By determining the H I content of the galaxies in the GEMS groups, we will be able to see if this trend continues in the less dense group environment. We have found two H I-deficient galaxies in the NGC 1566 group, NGC 1536 and 1515, which lie near or beyond the virial radius of the group. High-resolution H I observations of these two H I-deficient galaxies combined with H α measurements (e.g. Vogt et al. 2004) will provide further information to the past history of these galaxies, particularly whether their gas removal mechanism is due to tidal interactions, ram pressure stripping, or a combination of the two.

5 CONCLUSIONS

We have conducted a blind H I survey of a region of $5.5 \times 5.5 \text{ deg}^2$ in the NGC 1566 group. Thirteen galaxies were detected in the H I data cube, including two LSB dwarf galaxies that had no previously known redshifts (LSBG F157–081 and APMBGC 157+016+068), which are new group members. There are now 26 known members in the NGC 1566 group. Two H I detections have two or more optical counterparts, and are previously known interacting systems that are confused within our beam size of 15.5 arcmin (NGC 1533/IC 2038,

and NGC 1602/NGC 1596). No isolated H I clouds were detected in our survey to a limit of $\sim 3.5 \times 10^8 M_{\odot}$.

The total H I mass detected in the group was $3.5 \times 10^{10} M_{\odot}$, and the galaxy NGC 1566 itself contains nearly half of the H I mass of this group, with $M_{\text{HI}} = 1.4 \times 10^{10} M_{\odot}$.

We calculate the virial radius of this group to be 580 kpc, and find over half the galaxies we associate with the group to lie beyond this radius. Thus it is possible that this is a young, non-virialized group. The total H I contained in the late-type galaxies of this group appears to be consistent with the expected H I content based on their optical diameter and morphology. However there are two cases in which a spiral galaxy appears to be 5–10 times more deficient than expected: these are NGC 1536 and 1515. Further observations of these two galaxies are needed to determine their gas removal mechanisms.

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