TABLE I. ABSORPTION VELOCITIES IN KM/SEC, NOVA LACERTAE 1936

Date, U.T.	A	B	I	II	II'	III	IV	Orion	NIII
June 19.35		897	1110	1330					
20.24	571	886	1144	1316					
21.23	684	851	1356	1502	1639	2725	2956	2712	
22.21		870	1510	1687	1830	2862	3234	3184	
23.26		845	1638	1876	2089	3094	3347	3337	
24.20		912	1714	1997	2191	3184	3346	3336	
25.30	1205	904	1758	2096	2291	3272	3456	3362	
26.46		856	1804	2265			3358	3368	
27.38			1803	2336	2449	3108		3338	
28.30	1530	902	1920	2388		3171	3331	3197	
29.30	1598		1974	2428		3274*		2758	2878
30.37			1995	2475				3137	3516
July 1.34			2041	2484	2759	3410*		3396	3412
2.15			2100	2488	2722	3499*		2655	3286
3.23	1758		2080	2508	2685	3733*		3380	3492
4.22			2168	2514	2690	3427*		2790	
5.30	1746		2108	2539	2730	3296*		2680:	3152
6.22			2114	2527	2633	3409*		2748	2632
7.24			2098	2520	2649	3551*		3230	3248
8.22			2121	2509	2650			3648	3666
9.22			2130	2541				3440	3532
10.31			2097	2530		3617*		3499	3520
11.4			2114	2402		3535*		3792	3770
12.4				2453				4026†	3671

* $H\beta$ only.

† Nv lines.

Both Heard¹ and Harper² considered component A of the first few days to be the same as component I after July 1, and both measured blends of components I and II from June 21 to 24. Heard also connected component II of the first week with component II' after June 29, and component I of the first days with component II after June 29. Bertaud³ has followed Heard's interpretation, although the Mount Wilson spectrograms⁴ clearly show the continuity of component II from June 22 to July 7. Heard also connected components B and I of June 19 with I and II of June 20. Intensities of the two absorption systems and velocity changes during the night contradict such an interpretation.

The necessity of a continuous record and of the study of all available material by a single investigator is indicated. Compilations of data from several sources can easily lead to the perpetuation of error.

1. *Pub. A. A. S.* **9**, 42, 1937.2. *Pub. Dom. Astroph. Obs.*, Victoria **6**, 325, 1937.

3. Thesis, Paris, 1945.

4. *Pub. A. S. P.* **48**, 325, 1936.

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Mayall, Margaret W. Nova Sagittae 1913, another recurrent nova.

Another recurrent nova has been added to the small group of variable stars which seem to be intermediate between the Z Andromedae and U Geminorum or SS Cygni stars and the typical

novae. The list of recurrent novae now contains 6 stars, T CrB (1866, 1946), RS Oph (1898, 1933), T Pyx (1890, 1902, 1920), Nova Sge (1913, 1946), Nova Sgr (1901, 1919), and U Sco (1863, 1906, 1936).

Kurt Himpel, at the Heidelberg Observatory, in Germany, on the night of June 28-29, 1946, observed Nova Sge 1913 to increase from 12th to the 10th magnitude in 4 hours, or at the rate of $\frac{1}{2}$ magnitude per hour. A search of the photographs at the Harvard Observatory revealed a plate taken about 4 hours later than Himpel's second observation. The nova was then magnitude 8.9 ptg. The maximum observed, 7.7 mag., was reached on the next night, June 29-30. The brightness decreased steadily until it reached 13.3 on August 15, 48 days after maximum, the latest observation available.

Nova Sge 1913 was originally discovered by Miss J. Mackie, at the Harvard Observatory.¹ All the plates have been remeasured by the author, using a sequence determined by Miss Leavitt, and reduced to the international scale. Maximum occurred on November 22, 1913, magnitude 7.0 ptg. The nova faded steadily, but not quite so rapidly as at the 1946 maximum, and 35 days after maximum reached magnitude 10.4. The next season, the summer of 1914, it appeared to vary between 13.7 and 14.5. One plate in 1919 showed it 14.6 mag., and 2 plates in 1937 showed it 14.9 mag. on May 20 and 15.5 mag. on Sept. 1.

The parallax was measured by Van Maanen